New Insight on
Neutrino Dark Matter Interactions
from Small-Scale CMB Observations

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Annual Reporting Seminar 2023
Fundamental Research Department

December 13, 2023

TALK BASED ON ARTICLES

1)

Monthly Notices of the ROYAL ASTRONOMICAL SOCIETY

MNRAS **527,** L122–L126 (2024) https://doi.org/10.1093/mnrasl/slad157
Advance Access publication 2023 October 17

New insights on ν–DM interactions

Philippe Brax,[1,2]★ Carsten van de Bruck,[3]★ Eleonora Di Valentino,[3]★ William Giaré[3] and Sebastian Trojanowski[4,5]

2)

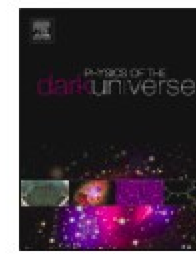
Physics of the Dark Universe 42 (2023) 101321

Contents lists available at ScienceDirect

Physics of the Dark Universe

journal homepage: www.elsevier.com/locate/dark

ELSEVIER



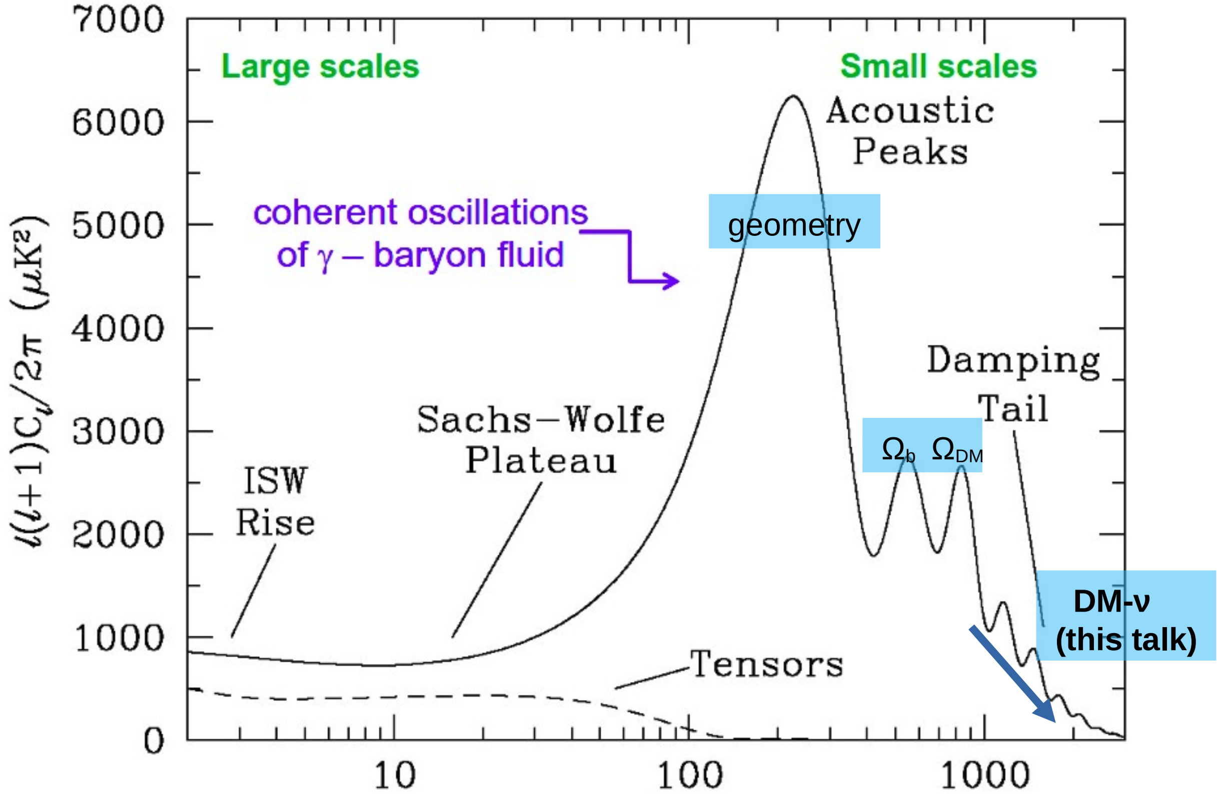
Extended analysis of neutrino-dark matter interactions with small-scale CMB experiments

Check for updates

CMB & DARK MATTER

Cosmic Microwave Background (CMB) radiation plays an essential role in establishing DM role in the evolution of the Universe

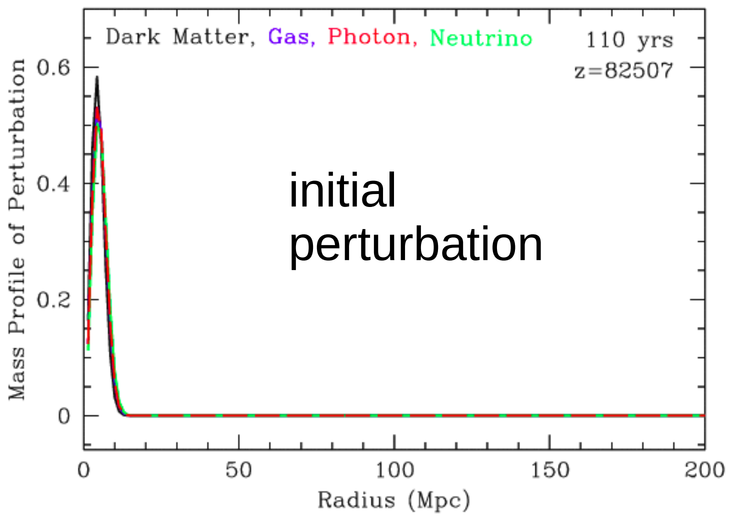
Credit: Carlos Frenk

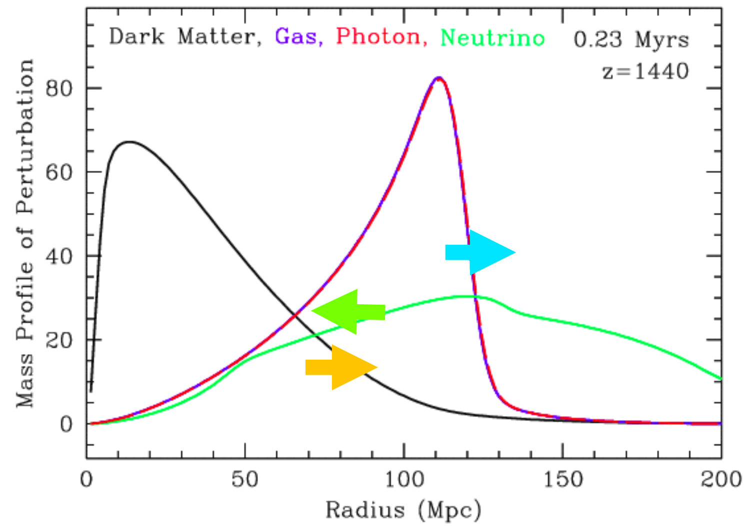


CAN WE LEARN MORE FROM THE DAMPING TAIL?

ACOUSTIC OSCILLATIONS & NEUTRINOS

Credit: Daniel Eisenstein





- Standard cosmology: neutrinos free stream & can "drag" baryon photon fluid

- In the presence of DM-neutrino interactions:

R.J. Wilkinson, etal, 1401.7597

– DM can take part in oscillations → gravitational boost & enhanced CMB peaks

G. Magano, etal 0606190

– DM-ν interactions can affect ν free streaming → stronger clustering & enhanced CMB peaks >1

– DM-neutrino fluid has a lower sound speed → drag effect, CMB peaks shifted
and more…

P. Serra, etal, 0911.4411

- **This talk:** low (but non-negligible) DM-ν interaction strength

$\Gamma_{\rm DM-\nu} > H > (\Gamma_\nu \equiv \Gamma_{\nu-e} + \Gamma_{\nu-\rm DM})$, **mixed damping regime**

C. Boehm, R. Schaeffer, 0410591
J. Stadler, C. Boehm, O. Mena, 1903.00540

SMALL-SCALE CMB & ν-DM INTERACTIONS

"Conventional wisdom" **WRONG !**

A. Paul, etal, 2104.04760

- The features of the tail of CMB PS is highly suppressed due to diffusion damping, so no visible difference due to non-standard scatterings is present at those modes. On

Our findings

$$u_{\nu\mathrm{DM}} = \frac{\sigma_{\nu\mathrm{DM}}}{\sigma_{\mathrm{T}}}\left(\frac{m_{\mathrm{DM}}}{100\ \mathrm{GeV}}\right)^{-1}$$

- DM-ν interactions:

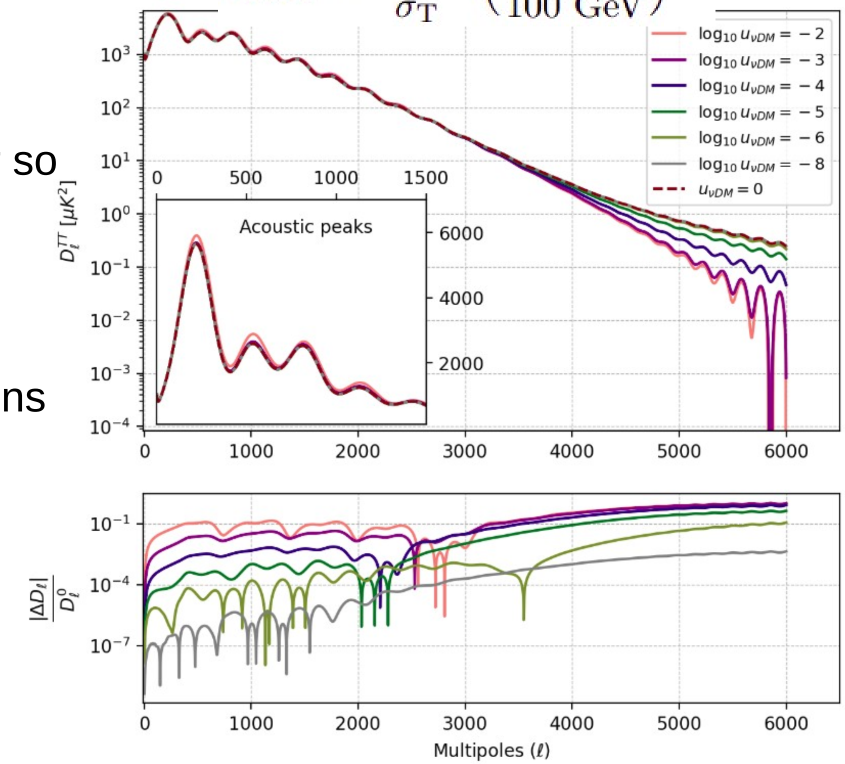
– suppression of high-multipole peaks at few % level or so

– negligible effect at low multipoles for $u_{\nu DM} < 10^{-5}$

- Similar effect in the temperature (TT) & polarization (EE) distributions

- Current data: Atacama Cosmology Telescope (ACT), South Pole Telescope (SPT)

- Future surveys can further improve: CMB-S4, ...



High-multipole CMB data = new window to study DM-ν interactions

PREFERENCE: NON-ZERO DM-ν COUPLING

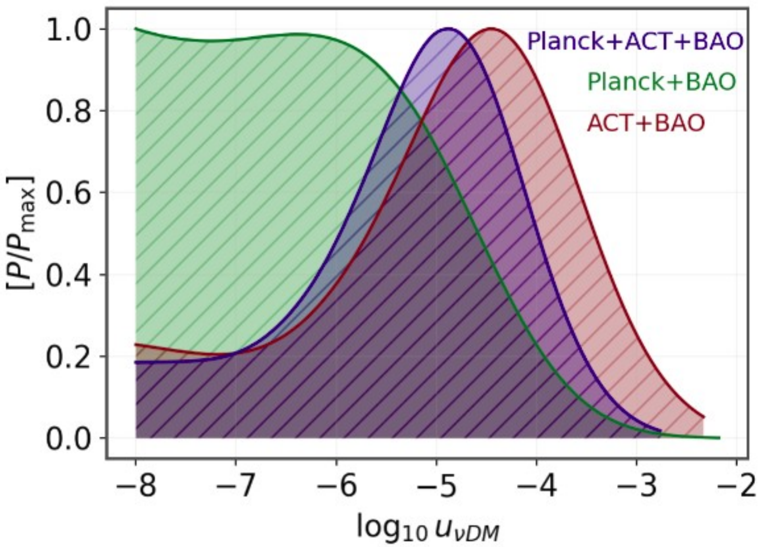
$\sigma_{\nu DM} \sim T^0$

CMB (this work)

Hints from Lyman-α

D.C. Hopper, M. Lucca, 2110.04024

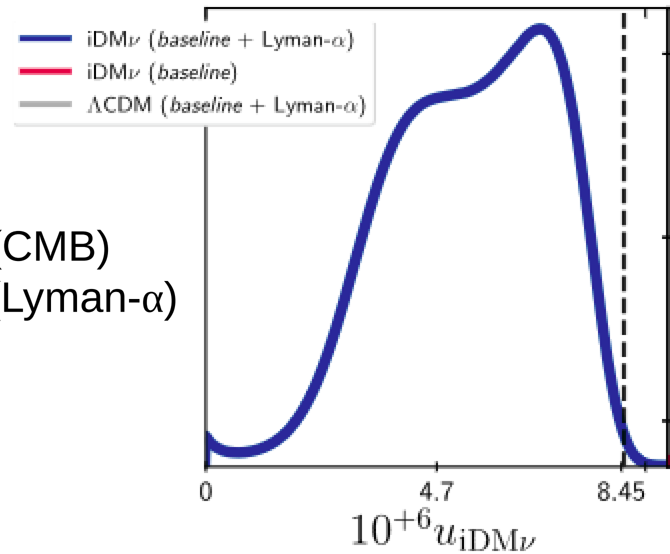
Posterior



$\mathrm{Log}_{10}\, u_{\nu DM} = -5.20^{+1.20}_{-0.74}$ (CMB)

$-5.42^{+0.17}_{-0.08}$ (Lyman-α)

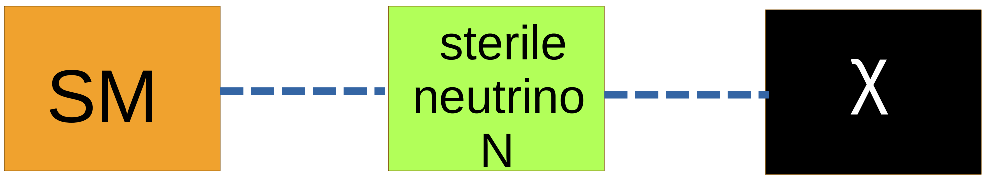
$$u_{\nu DM} = \frac{\sigma_{\nu DM}}{\sigma_T}\left(\frac{m_{DM}}{100\ \mathrm{GeV}}\right)^{-1}$$

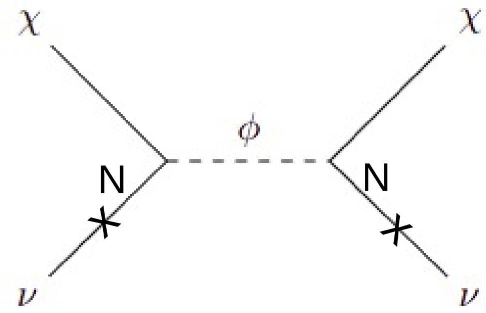


Parameter	Planck	Planck + BAO	ACT	ACT + BAO	ACT + Planck + BAO
$\Omega_b h^2$	0.02239 ± 0.00015	0.02239 ± 0.00013	0.02153 ± 0.00030	0.02154 ± 0.00030	0.02236 ± 0.00012
$\Omega_c^{\nu DM} h^2$	0.1196 ± 0.0012	0.11958 ± 0.00093	0.1185 ± 0.0039	0.1198 ± 0.0015	0.11975 ± 0.00097
$100\theta_s$	1.04193 ± 0.00030	1.04191 ± 0.00028	1.04337 ± 0.00069	1.04321 ± 0.00063	1.04206 ± 0.00026
τ_{reio}	0.0528 ± 0.0074	0.0524 ± 0.0072	0.064 ± 0.015	0.062 ± 0.014	0.0563 ± 0.0064
$\log(10^{10} A_s)$	3.039 ± 0.014	3.038 ± 0.014	3.049 ± 0.030	3.047 ± 0.030	3.053 ± 0.013
n_s	0.9642 ± 0.0044	0.9642 ± 0.0038	1.004 ± 0.016	1.001 ± 0.014	0.9678 ± 0.0036
$log_{10} u_{\nu DM}$	$< -4.42\ (< -3.95)$	$< -4.46\ (< -4.39)$	$-5.08^{+1.5}_{-0.98}\ (< -3.74)$	$-4.86^{+1.5}_{-0.83}\ (< -3.70)$	$-5.20^{+1.2}_{-0.74}\ (< -4.17)$
H_0	$68.03 \pm 0.55\ (68.0^{+1.1}_{-1.1})$	$68.05 \pm 0.42\ (68.05^{+0.81}_{-0.82})$	$68.2 \pm 1.6\ (68.2^{+3.3}_{-3.3})$	$67.66 \pm 0.58\ (67.7^{+1.1}_{-1.2})$	$68.01 \pm 0.43\ (68.01^{+0.83}_{-0.85})$
σ_8	$0.806^{+0.013}_{-0.0097}\ (0.806^{+0.024}_{-0.028})$	$0.807^{+0.011}_{-0.0084}\ (0.807^{+0.020}_{-0.021})$	$0.823^{+0.025}_{-0.021}\ (0.823^{+0.046}_{-0.050})$	$0.821^{+0.025}_{-0.020}\ (0.821^{+0.044}_{-0.050})$	$0.820^{+0.011}_{-0.0093}\ (0.820^{+0.021}_{-0.023})$
$\ln BF$	-3.74	-2.48	-0.194	-0.156	0.525

B. Bertoni, etal, 1412.3113
B. Batell, etal, 1709.07001
P. Brax, etal, 2305.01383

STERILE NEUTRINO PORTAL TO DARK MATTER





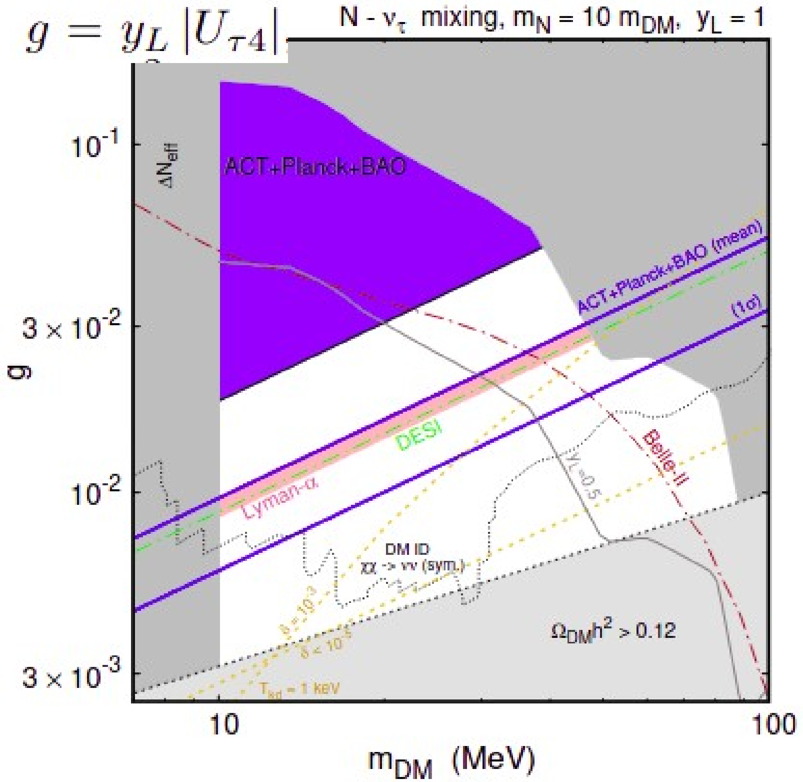
$$\mathcal{L} \supset -\phi\,\bar{\chi}\,(y_L\,N_L + y_R\,N_R) + \text{h.c.}$$

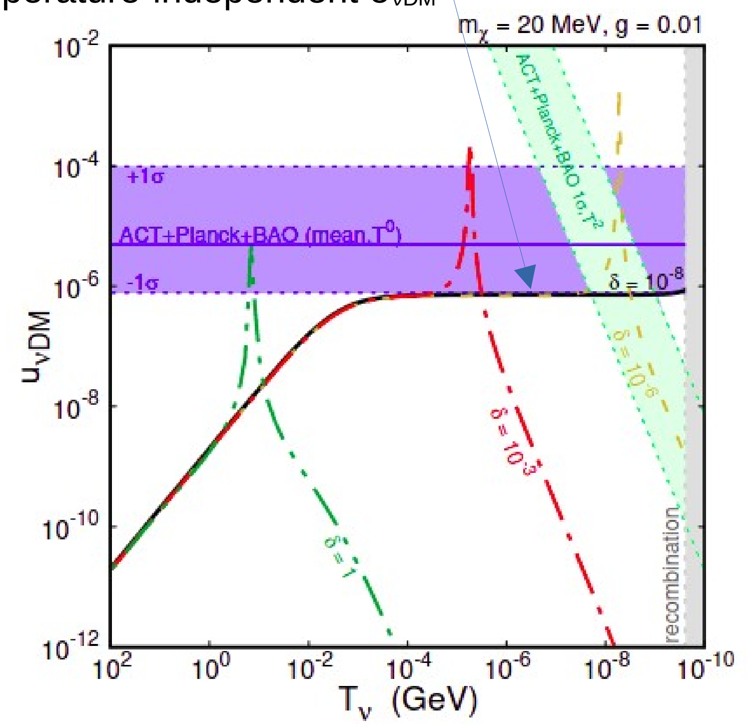
$$\mathcal{L} \supset -\lambda(\bar{L}\,\tilde{H})\,N_R,$$

$$\delta = (m_\phi - m_\chi)/m_\chi. << 1$$

temperature-independent $\sigma_{\nu DM}$

$$g = y_L\,|U_{\tau 4}|$$





CONCLUSIONS

- CMB observations are crucial for our understanding of dark matter

- **small-scale CMB measurements with few % accuracy open a new window to study DM interactions with neutrinos**

- **preference for non-zero DM-ν coupling in the high-multipole ACT data & agreement with low-multipole Planck data + BAO & RSD**

- Similar earlier hints from Lyman-α

- Toy model: sterile neutrino portal to DM

- Can accommodate the data but careful checking of other effects needed (cutoff scale, DM self-interactions...)

- Future data: ACT, CMB-S4, DESI, … + accelerator-based bounds on sterile neutrinos

THANK YOU !

BACKUP

CURRENT DATA & ANALYSIS

DATA

- **Planck** 2018

temperature & polarization
1907.12875, 1807.06209, 1807.06205

lensing 1807.06210

- Atacama Cosmology Telescope (**ACT**)

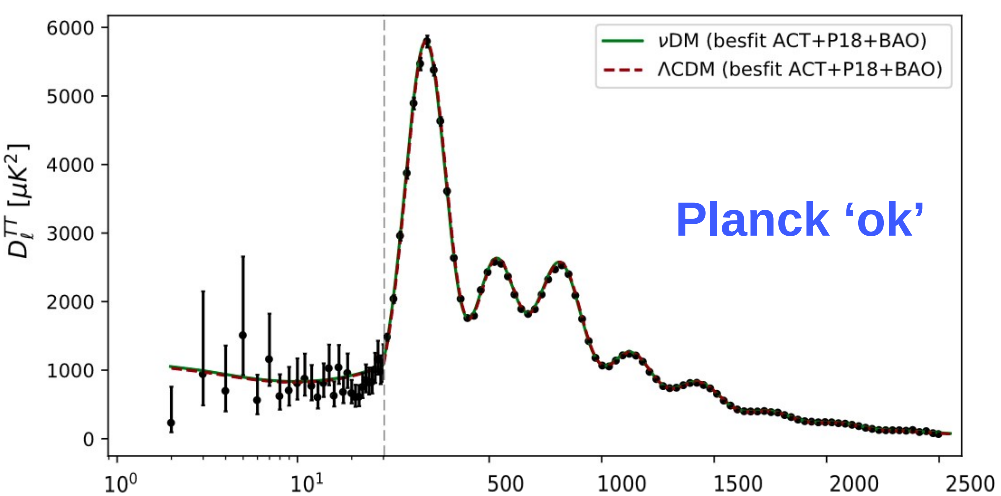
temp. & polar. DR4 2007.07289

new confirmation: + ACT-DR6 & SPT W. Giare, etal, 2311.09116

- Baryon Acoustic Oscillations (**BAO**)

& Redshift Space Distortions

BOSS DR12 1208.0022



ANALYSIS

- (modified) CLASS + DM-ν

1104.2933, 1903.00540, 2011.04206

- Sampling:

COBAYA (with CosmoMC)
2005.05290, 0205436, 1304.4473

Parameter	$\sigma_{\nu\mathrm{DM}} \sim T^0$
$\Omega_b h^2$	$[0.005\,,\,0.1]$
$\Omega_c^{\nu\mathrm{DM}} h^2$	$[0.005\,,\,0.1]$
$100\,\theta_{\mathrm{MC}}$	$[0.5\,,\,10]$
τ	$[0.01\,,\,0.8]$
$\log(10^{10} A_S)$	$[1.61\,,\,3.91]$
n_s	$[0.8\,,\,1.2]$
N_{eff}	$[0\,,\,10]$
DM-ν $\log_{10} u_{\nu\mathrm{DM}}$	$[-8\,,\,-1]$

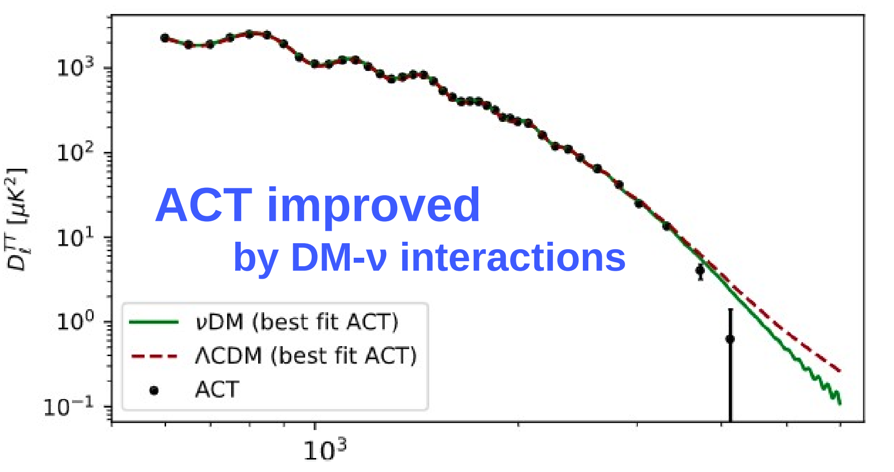
$$u_{\nu\mathrm{DM}} = \frac{\sigma_{\nu\mathrm{DM}}}{\sigma_T}\left(\frac{m_{\mathrm{DM}}}{100\ \mathrm{GeV}}\right)^{-1}$$

σ_T – Thomson scat.

- Adding ACT:

- weaker bounds on $u_{\nu\mathrm{DM}}$

- non-zero coupling preferred



B. Bertoni, etal, 1412.3113
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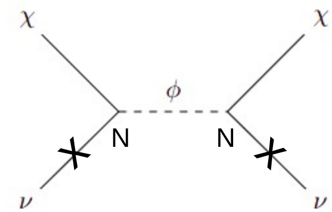
P. Brax, etal, 2305.01383

SAMPLE MODEL – STERILE NEUTRINO PORTAL

- Fermionic χ DM coupled to a new scalar ϕ and (heavy) sterile neutrino N

$$\mathcal{L} \supset -\phi\,\bar{\chi}\,(y_L\,N_L + y_R\,N_R) + \text{h.c.}$$



- Mixing with the active neutrino (dominant with ν_τ)

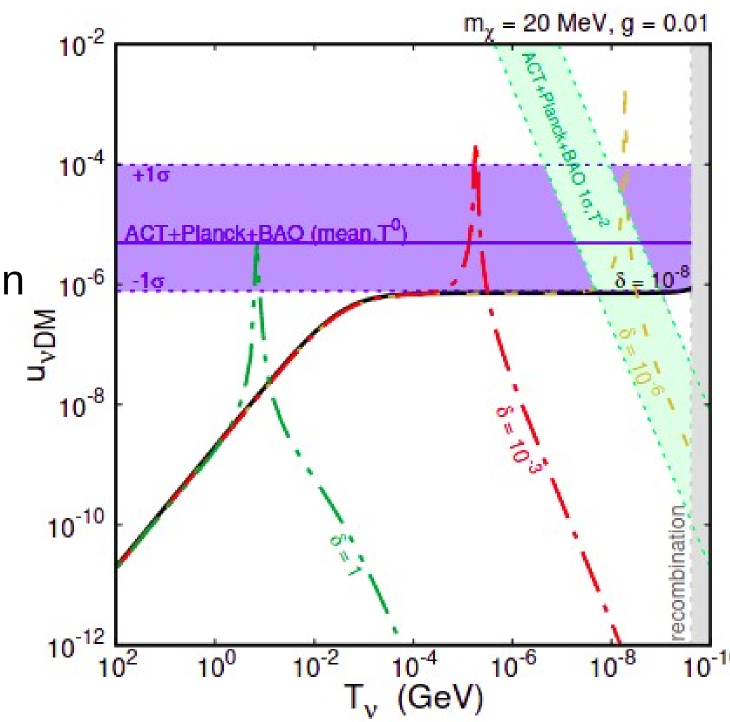
$$\mathcal{L} \supset -\lambda\,(\bar{L}\,\tilde{H})\,N_R,$$

- **For small mass splittings DM-ν cross section ~ T^0**

$$\sigma_{\chi\nu} \simeq (10^{-34}\,\text{cm}^2)\left(\frac{g}{0.01}\right)^4\left(\frac{20\text{ MeV}}{m_\chi}\right)^2 \times\left[1+0.075\left(\frac{m_\chi}{20\text{ MeV}}\right)^2\left(\frac{T_{\text{rec.}}}{T_\nu}\right)^2\left(\frac{\delta}{10^{-8}}\right)^2\right]$$

$$\delta = (m_\phi - m_\chi)/m_\chi << 1$$

for a range of temperatures



- For lower temperatures, σ grows, resonance ϕ production

- At even lower temperatures, $\sigma \sim T^2 / m_\chi^4$

$$\sigma_{\chi\nu} \simeq (10^{-52}\,\text{cm}^2)\left(\frac{g}{0.1}\right)^4\left(\frac{100\text{ MeV}}{m_\phi}\right)^4\left(\frac{T}{T_0}\right)^2$$

- At high temperatures, $\sigma \sim 1/T^2$

– helps avoiding astrophysical bounds from blazars, etc.

FITTING THE DATA

● The model can accommodate the data
& avoid exclusion bounds from accelerator-based searches

● DM relic density requires, e.g., asymmetric DM component

● This also helps avoiding DM ID bounds from χχ → νν

● Future probes: DESI, Belle-II (τ decays)

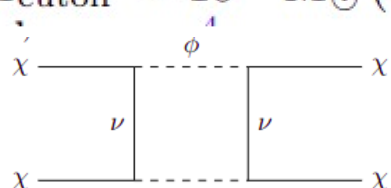
CHALLENGES OF THE TOY MODEL:

● Mass degeneracy requires fine-tuning
& can be radiatively unstable via $\lambda_{\phi H}|\phi|^2|H|^2$

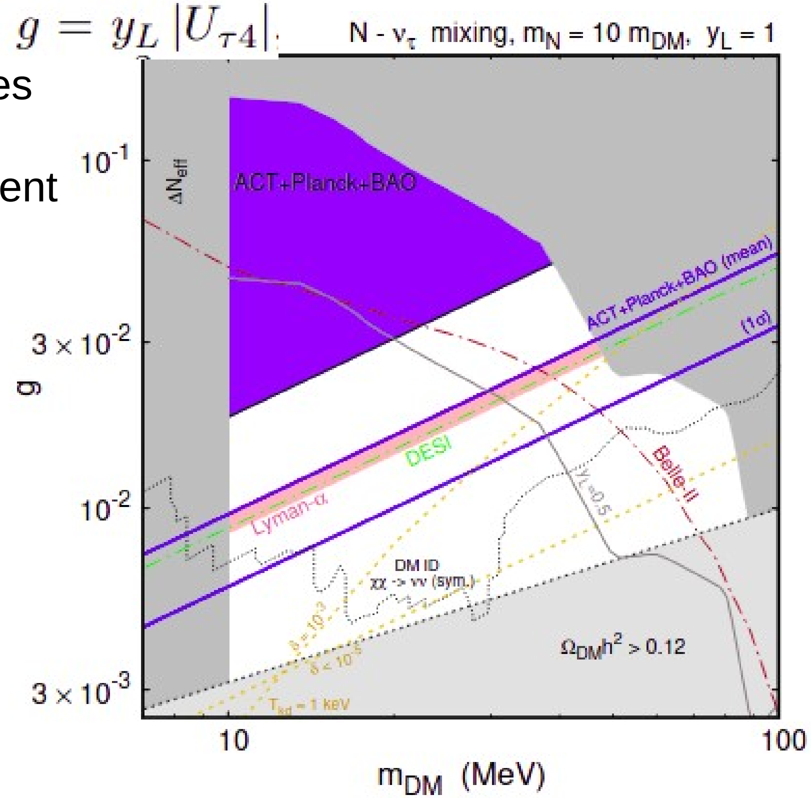
● Low neutrino kinetic decoupling temperature

$$T_{kd}\Big|_{m_\phi \simeq m_\chi} \simeq (0.12\ \mathrm{keV}) \left(\frac{0.01}{g}\right)^2 \left(\frac{m_\chi}{20\ \mathrm{MeV}}\right)^{3/2} \longrightarrow M_{\rm cutoff} \sim 10^{11}\, M_\odot\, (0.1\ \mathrm{keV}/T_{kd})^3$$

● Possible DM (too) strong self-interactions



$g = y_L\,|U_{\tau 4}|$



POSSIBLE SOLUTIONS: fraction of DM interacts strongly with νs; increase the mass splitting δ